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**THE STUDY OF LARGE SCALE STRUCTURE**

Grant NAGW-201

**FINAL REPORT**

**For the period 01 July 1981 through 30 June 1996**

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During 1995-6, NAGW-201 provided a small amount of support for our continuing studies of systems of galaxies and large-scale structure (see bibliography below).

The Century Survey is now complete to a *well-calibrated*, somewhat fainter limiting apparent magnitude than we had originally planned ( $m_R = 16.1$ ). Perhaps the most striking result of the survey is the evidence a feature in the power spectrum at a scale of  $\sim 5000 \text{ km s}^{-1}$ ; i.e. a “preferred” scale for the voids. From a physical point of view, this scale is particularly interesting: it is approximately the scale of the horizon at the epoch of decoupling.

We have computed a preliminary luminosity function, the spatial correlation function, and the velocity moments. The velocity moments as a function of scale are in striking agreement with the CfA+SSRS2 results. Although the Century Survey contains one-tenth the number of galaxies in the CfA+SSRS2, it samples a comparable number of large ( $\sim 5000 \text{ km/s}$ ) structures. For this reason (along with the better galaxy positions and well-calibrated magnitudes), the velocity moments are well constrained. We plan to publish analyses and the data for the Century Survey within the next year.

To investigate some of the scientific issues raised by analyses of the CfA, SSRS2, and Century surveys (see e.g. Marzke *et al.* 1995), we have nearly completed a survey complete to  $m_R = 15.4$  (15R survey) over  $\sim 1000$  square degrees. The striking preliminary maps again show well-defined scale of  $\sim 5,000 \text{ km/s}$  for the largest voids. The survey also indicates that, as suggested by recent analyses of the LCRS survey, there is modulation of the structure on a scale of  $\gtrsim 100 \text{ Mpc}$ . We plan to submit a first paper this fall.

We have nearly completed the process of obtaining arcsecond coordinates for all of the galaxies in the CfA redshift survey. These coordinates will enable the *first* computation of the pairwise velocity dispersion for the galaxy distribution (Marzke *et al.* 1995) on scales of  $100 \text{ kpc}$  and smaller. The small-scale pairwise dispersion is a measure of the importance of (1) the dynamical state of systems and (2) the “softness” of galaxy-galaxy interactions (Diaferio and Geller 1996). This new study complements Barton’s recently completed work on selection of compact groups from redshift surveys (Barton *et al.* 1996).

Direct distance measurements enable more direct studies of large-scale dynamics. We have completed an IRTF investigation of the velocity field in the vicinity of the Great Wall (Dell’Antonio *et al.* 1996a, b). We derived the first limits on the spatial thickness of the Wall ( $\lesssim 11h^{-1} \text{ Mpc}$ ) and thus demonstrate that the Great Wall is indeed a thin 2D structure in real as well as in redshift space. We also show that there is minimal shear across the extent of the Great Wall. We plan to continue this project when the Arecibo upgrade is completed this fall; we should be able substantially better physical constraints on the properties of the Great Wall and on large-scale flows on the  $10,000 \text{ km/s}$  scale.

We also continued to study groups (Ramella *et al.* 1995 a,b; Diaferio *et al.* 1995;

Dell'Antonio *et al.* 1996) and clusters of galaxies (Mahdavi *et al.* 1996). Mohr *et al.* (1995) extendend their examination of the cluster morphology-cosmology connection. They are in the process of investigating how optical data improve the constraints imposed by x-ray observations (Mohr *et al.* 1996a, b).

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